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COSMOLOGICAL EFFECTS OF THE FIRST STARS: EVOLVING SPECTRA OF POPULATION III

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ABSTRACT

The first stars hold intrinsic interest for their uniqueness and for their potentially important contributions to galaxy formation, chemical enrichment, and feedback on the intergalactic medium (IGM). Although the sources of cosmological reionization are unknown at present, the declining population of large bright quasars at redshifts $z > 3$ implies that stars are the leading candidates for the sources that reionized the hydrogen in the IGM by $z \sim 6$. The metal-free composition of the first stars restricts the stellar energy source to proton-proton burning rather than the more efficient CNO cycle. Consequently, they are hotter, smaller, and have harder spectra than their present-day counterparts of finite metallicity. We present new results from a continuing study of metal-free stars from a cosmological point of view. We have calculated evolving spectra of Population III clusters, derived from a grid of zero-metallicity stellar evolutionary tracks. We find that H-ionizing photon production from metal-free stellar clusters takes twice as long as that of Population II to decline to 1/10 its peak value. In addition, metal-free stars produce substantially more photons than Population II in the He II ($E > 4$ ryd) continuum. We suggest that large Ly α equivalent widths ($W_{\text{Ly}\alpha} > 400 \text{ \AA}$) may provide a means of detecting metal-free stellar populations at high redshift and that He II recombination lines ($\lambda\lambda 1640, 4686$) may confirm identifications of Population III. While Population III clusters are intrinsically bluer than their Population II counterparts, nebular continuum emission makes up this difference and may confuse attempts to discern Population III stars with broadband colors. In a companion paper, we explore the consequences of evolving spectra of Population III for the reionization of the IGM in both H and He.

Subject headings: cosmology: theory — galaxies: high-redshift — stars: evolution

1. INTRODUCTION

The first stars and their role in galaxy formation and the evolution of the intergalactic medium (IGM) define one of the frontiers of modern cosmology, yet we do not know when the first stars formed, where they now reside, what happened to them over time, or even where to look for them. The big bang paradigm requires that the first stars be composed of H and He with only trace light elements and negligible carbon (Schramm & Turner 1998). Metal-free stars have unique properties owing to this primordial composition: they are hotter and smaller than their metal-enriched counterparts (Tumlinson & Shull 2000, hereafter TS). These features have interesting implications for the formation, environment, and particularly the spectra of the first stars. Some key open issues are the formation epoch of the first stars (Ricotti, Gnedin, & Shull 2002), the mass distribution for stars formed from primeval gas (Bromm, Coppi, & Larson 1999), and their contribution to the reionization of H and He in the IGM. The recent suggestions by Lanzetta et al. (2002) that star formation may have been more widespread in the distant past, the discovery of a gravitationally lensed galaxy at $z = 6.56$ (Hu et al. 2002), and the apparent completion of H reionization by $z \simeq 6$ (Becker et al. 2001; Djorgovski et al. 2001) and He at $z \simeq 3$ (Kriss

et al. 2001) make detailed examination of the first stars particularly timely.

With our present tools we can see back to $z \sim 6$, roughly 1 Gyr after the big bang (a universe with $\Omega_m = 0.3$, $\Omega_\Lambda = 0.7$, and $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$ is assumed here). Imaging and spectroscopic surveys that sample wide swaths of sky for telltale emission lines (Ly α emitters; Hu, McMahon, & Cowie 1999; Rhoads et al. 2000) or peculiar broadband colors (Lyman break galaxies; Steidel et al. 1999) have discovered more than 1000 galaxies at $z \sim 3$. Searches such as the Sloan Digital Sky Survey are expanding the list of known QSOs at $z \gtrsim 5$ and enabling high-redshift probes of the IGM (Fan et al. 2002; Schneider et al. 2002). Ground-based spectra can determine the star formation rate (SFR), metallicity, and mass of galaxies at $z = 3$ (Pettini et al. 2002). The first stars are believed to have formed between $z = 20$ and 30 (Gnedin & Ostriker 1997; Ricotti et al. 2002), while galaxies have been discovered up to $z \sim 6$. The next observational frontier is the direct observation of the first stellar generation.

This work is part of an ongoing effort to examine the first stars from a unified stellar and cosmological point of view. We began with zero-age models (TS) of metal-free stars that revealed their potential importance to H and He reionization in the IGM. We followed that with a detailed calculation of the He II recombination-line fluxes from Population III clusters, assessing the prospects for discovering this emission as a distinctive signature of the first stars (Tumlinson, Giroux, & Shull 2001, hereafter TGS). We are concerned here with the properties of the first stars that affect their observational signatures and feedback on the IGM.

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These effects depend on the answer to a fundamental question about the first stars: What is the time evolution of their structure and radiation? We update the work of TS to compute new evolving spectra of metal-free stars and to assess the effects of evolution on the emergent spectra of Population III clusters.

In this paper, we use the terms “Population III” and “metal-free” interchangeably. We also note that there are two distinct usages of the term “Population III” that have appeared in the literature. One usage connotes small protogalactic clusters that have $M_{\text{DM}} \simeq 10^6\text{--}10^7 M_{\odot}$ and that must cool with molecular hydrogen lines in the absence of metals. We adhere to the stellar definition, in which “Population III” and metal-free stars are identical, regardless of location.

In § 2 we discuss the major components of our stellar evolution calculations. In § 3 we present the stellar evolution tracks that enable our evolving spectra and observational predictions. In § 4 we describe the modeling of the stellar atmospheres and construct evolving spectra for synthetic Population III clusters. In § 5 we make detailed predictions of the observational signatures of metal-free stars and describe the realistic prospects for their discovery. In § 6 we draw some general conclusions. Paper II (Venkatesan, Tumlinson, & Shull 2003) presents the results of a cosmological reionization model that incorporates Population III spectra as ionizing sources.

2. STELLAR EVOLUTION AT ZERO METALLICITY

2.1. Method

To study the evolution of Population III stars, we must construct realistic numerical models that account for the physical processes that are thought to be relevant in real stars. The basic theory of stellar structure is quite well developed and has been tested against observations for decades. Occasional conflicts with detailed observations lead to refinements of the theory, usually in the form of additional input physics (convective overshoot, mass loss, and rotation, for example). Of course, no such fine calibration of models for Population III stars is yet possible, since we have not discovered any of these stars or anything believed to be their remnants. We must make the basic assumption that the framework of stellar evolution theory applies equally well to the unusual regime of zero metallicity. Under this assumption, our Population III models incorporate the same paradigmatic assumptions and physical ingredients that underlie all modern stellar evolution codes. We must note here, however, that since our ultimate goal is a calculation of the evolving spectra, we are not concerned with internal processes, particularly late in the stellar lifetime, that do not affect the spectral evolution. Those readers who are interested in the details of the later stages of Population III evolution are referred to the tracks by Marigo et al. (2001) and Siess, Livio, & Lattanzio (2002). Schaerer (2002) has compiled tracks from a variety of sources (El Eid, Fricke, & Ober 1983; Marigo et al. 2001; Klapp 1983) in an investigation into the properties of metal-free stars that parallels some of our own results. Finally, we note that the following exposition of our method is included for readers interested in the details. The basic results of our study can be found by skipping to § 3.

In our code we solve the fundamental stellar structure equations:

$$\frac{dR}{dM_r} = \frac{1}{4\pi R^2 \rho}, \quad (1)$$

$$\frac{dL_r}{dM_r} = \epsilon - T \frac{dS}{dt}, \quad (2)$$

$$\frac{dP}{dM_r} = -\frac{GM_r}{4\pi R^4}, \quad (3)$$

$$\frac{dT}{dM_r} = -\frac{GM_r T}{4\pi R^4 P} \nabla, \quad (4)$$

where radius R , luminosity L_r , pressure P , and temperature T are the “primary variables” and the subscript r implies that the variable represents a cumulative quantity up to radius R . These equations are transformed to a Lagrangian mass grid that does not change with total stellar mass. The mass density ρ , opacity κ , energy generation rate ϵ , thermodynamic derivatives ∇_{rad} and ∇_{ad} , and specific entropy S are “secondary variables” that are themselves functions of T , P , L , R , and the chemical abundances by number, denoted by X_i . For radiative energy transport, we use $\nabla = \nabla_{\text{rad}} = (d \ln T / d \ln P)_{\text{rad}}$, and in convective regions we use $\nabla = (d \ln T / d \ln P)_{\text{ad}}$. The adiabatic derivative is calculated for each mass point during the evaluation of the equation of state.

We solve these coupled equations with a variant of the relaxation method described definitively by Kippenhahn, Weigert, & Hofmeister (1967) and still in extensive use today. This method transforms the main *differential* equations into *difference* equations by expanding them in a Taylor series about small corrections to the initial guesses for the primary variables. The corrections are derived iteratively by a routine adapted from Press et al. (1987, chap. 16.3).

2.2. Physical Ingredients of the Models

2.2.1. Equation of State

The equation of state (EOS) is a numerical function that relates the primary variables pressure P and temperature T and the secondary variable density ρ . The EOS is that of an ideal monatomic gas with radiation pressure. Partial ionization of the metallic species is included but is generally unimportant at $T > 10^7$ K in the stellar core.

2.2.2. Opacity

Our code computes the Rosseland mean opacity $\kappa = \kappa(\rho, T, X, Z)$ from the OPAL tables produced at the Lawrence Livermore National Laboratory (Iglesias & Rogers 1996). The code calls a subroutine provided publicly by the OPAL project to interpolate in the tables and derive κ for a specific combination of parameters (ρ, T, X, Z) . The tables are provided at fixed grid points in T and R [$R = \rho / (T / 10^6 \text{ K})^3$, not to be confused with the radius], and there are multiple tables corresponding to different metallicities.

2.2.3. Energy Generation

Energy generation rates (ϵ ; in $\text{ergs s}^{-1} \text{g}^{-1}$) are computed by a custom nuclear reaction network. The network stores the relative populations of the relevant nuclear species and derives from these and the reaction rates the energy

TABLE 1
REACTIONS INCLUDED IN THE STELLAR
EVOLUTION MODELS

Class	Reaction
<i>p-p</i> I	${}^1\text{H} + {}^1\text{H} \rightarrow {}^2\text{H} + \nu_e + e^+$
	${}^2\text{H} + {}^1\text{H} \rightarrow {}^3\text{He} + \gamma$
	${}^3\text{He} + {}^3\text{He} \rightarrow {}^4\text{He} + 2 {}^1\text{H}$
<i>p-p</i> II	${}^3\text{He} + {}^4\text{He} \rightarrow {}^7\text{Be} + \gamma$
	${}^7\text{Be} + e^- \rightarrow {}^7\text{Li} + \nu_e$
	${}^7\text{Li} + {}^1\text{H} \rightarrow 2 {}^4\text{He}$
<i>p-p</i> III.....	${}^7\text{Be} + {}^1\text{H} \rightarrow {}^8\text{B} + \gamma$
	${}^8\text{B} \rightarrow {}^8\text{Be}^* + e^+ + \nu_e$
	${}^8\text{Be}^* \rightarrow 2 {}^4\text{He}$
CNO.....	${}^{12}\text{C} + {}^1\text{H} \rightarrow {}^{13}\text{N} + \gamma$
	${}^{13}\text{N} \rightarrow {}^{13}\text{C} + e^+ + \nu_e$
	${}^{13}\text{C} + {}^1\text{H} \rightarrow {}^{14}\text{N} + \gamma$
	${}^{14}\text{N} + {}^1\text{H} \rightarrow {}^{15}\text{O} + \gamma$
	${}^{15}\text{O} \rightarrow {}^{15}\text{N} + e^+ + \nu_e$
	${}^{15}\text{N} + {}^1\text{H} \rightarrow {}^{12}\text{C} + {}^4\text{He}$
	${}^{15}\text{N} + {}^1\text{H} \rightarrow {}^{16}\text{O} + \gamma$
	${}^{16}\text{O} + {}^1\text{H} \rightarrow {}^{17}\text{F} + \gamma$
	${}^{17}\text{F} \rightarrow {}^{17}\text{O} + e^+ + \nu_e$
	${}^{17}\text{O} + {}^1\text{H} \rightarrow {}^{14}\text{N} + {}^4\text{He}$
	${}^{17}\text{O} + {}^1\text{H} \rightarrow {}^{18}\text{O} + e^+ + \nu_e$
Triple- α	${}^4\text{He} + {}^4\text{He} \rightarrow {}^8\text{Be}^*$
	${}^8\text{Be}^* \rightarrow 2 {}^4\text{He}$
α -capture	${}^8\text{Be}^* + {}^4\text{He} \rightarrow {}^{12}\text{C} + \gamma$
	${}^{12}\text{C} + {}^4\text{He} \rightarrow {}^{16}\text{O} + \gamma$
	${}^{16}\text{O} + {}^4\text{He} \rightarrow {}^{20}\text{Ne} + \gamma$
	${}^{20}\text{Ne} + {}^4\text{He} \rightarrow {}^{24}\text{Mg} + \gamma$
	${}^{14}\text{N} + {}^4\text{He} \rightarrow {}^{18}\text{F} + \gamma$
	${}^{18}\text{O} + {}^4\text{He} \rightarrow {}^{22}\text{Ne} + \gamma$
	${}^{22}\text{Ne} + {}^4\text{He} \rightarrow {}^{25}\text{Mg} + n$

generation rates from proton-proton, CNO, and He burning at supplied values of the primary and secondary variables. The 24 isotopic abundances appearing in Table 1 are followed individually in a fully implicit, time-dependent network modeled after the formalism in Hix & Thielemann (1999). Table 1 also lists the included reactions, for which the reaction rates are drawn from the canonical Caughlan & Fowler (1988) compilation.

The nuclear reaction network is coupled to the stellar structure equations through the secondary variable ϵ and the abundances of the chemical species. These are stored in a separate array, and the hydrogen, helium, and metal mass fractions X , Y , and Z , respectively, are computed from this array, with scaling to the total density for use in the EOS and opacity routines.

2.2.4. Convection and Mixing

Where the temperature gradient is too steep for radiative energy transport, convective motions must carry the energy produced in the stellar core. We assume that the isotopic abundances are fully mixed in convective regions and that convective regions overshoot their formally defined boundaries. We follow the common practice (Maeder & Meynet 1987; Schaller et al. 1992) that assumes an overshooting distance of $0.20H_p$, where H_p is the pressure scale height evaluated at the convective boundary:

$$H_p = P \left| \frac{dR}{dP} \right| = \frac{P}{(GM_r/r^2)\rho}. \quad (5)$$

Different assumptions for the overshooting distance may contribute modestly to the systematic uncertainty in the behavior of the stars in the H-R diagram and their emergent spectra.

2.2.5. Boundary Conditions

Four boundary conditions are required to complete the system of four equations (eqs. [1]–[4]). The outer boundary conditions constrain the temperature and pressure at the photosphere of the star. We assume that the temperature at the photosphere is the effective temperature given by the relation $L = 4\pi R_*^2 \sigma_{\text{SB}} T_{\text{eff}}^4$.

To calculate the atmospheric pressure, we use a function that derives $P_{\text{atm}} = P(T, \rho, X, Y, Z)$ for the stellar photosphere. Near the stellar surface, we neglect changes in mass, radius, and luminosity and assume that $M_r = M_*$, $R = R_*$, and $L_r = L_*$. This region obeys a simple relationship common to all stratified atmospheres: $d\tau = -\kappa\rho dr$. Using this relation and combining equations (1) and (3), we derive a differential equation relating optical depth τ and pressure P :

$$\frac{dP}{d\tau} = \frac{GM_*}{\kappa R_*^2}. \quad (6)$$

Then, using the Eddington approximation to radiative transport, we write the relation

$$T^4 = \frac{3}{4} \left(\frac{L_*}{4\pi\sigma_{\text{SB}}R_*^2} \right) \left(\tau + \frac{2}{3} \right). \quad (7)$$

The boundary condition for the differential equation (6) is that the gas pressure vanishes at $\tau = 0$, leaving $P = P_{\text{rad}} = \frac{1}{3}(4\sigma_{\text{SB}}/c)T^4$. At the photosphere, $T = T_{\text{eff}}$, and solving for τ in equation (7) gives $\tau = \frac{2}{3}$. We then integrate equation (6) from $\tau = 0$ to $\tau = \frac{2}{3}$ and find the value of P corresponding to T_{eff} at the photosphere. This calculation is done at each iteration in the approach to convergence.

At the inner boundary we use the so-called central expansions (Kippenhahn et al. 1967), which fix the luminosity and radius at the inner mass zone.

2.2.6. Initial Conditions

The starting model for an evolutionary sequence is calculated at the zero-age main sequence (ZAMS). A mass and chemical composition are specified, and a static structure model is calculated at $t = 0$. In this case, the starting models were the ZAMS models presented by TS. We use the initial composition $X = 0.76$, $Y = 0.24$ to represent the primordial composition. For illustration, we note that a decrease in X leads to systematically higher core temperatures for stars powered by *p-p* and CNO burning ($\epsilon_{p-p} \propto X^2$ and $\epsilon_{\text{CNO}} \propto XZ_C$).

For $M \gtrsim 15 M_{\odot}$, ZAMS Population III stars are not strictly metal-free except perhaps in the early stages of their formation (El Eid et al. 1983; Castellani, Chieffi, & Tornambe 1983; Marigo et al. 2001). The stellar cores are hot enough to accumulate a small abundance of ${}^{12}\text{C}$ via the triple- α process before and during main-sequence (MS) H burning. Thus, the idea of a high-mass, metal-free, MS star powered solely by *p-p* burning is misleading. In fact, these stars derive most of their energy from CNO burning of H (see § 3.2 below). We must carefully follow the production of ${}^{12}\text{C}$ in the core to derive the initial conditions for an MS Population III model. In contrast, TS used ad hoc

assumptions for the pre-MS synthesis of carbon, based on the earlier calculations of El Eid et al. (1983).

We start with the strictly metal-free ZAMS model from TS and let it produce ^{12}C in the core until CNO burning dominates the energy production. Starting from the $Z = 0$ model, the star evolves to slightly lower luminosity and T_{eff} until energy production from CNO burning exceeds that of the p - p chains. At this point, the star begins to evolve to higher luminosity but still declines in T_{eff} . The ^{12}C mass fraction at this transition is $Z_C \sim 10^{-12}$ (where $Z_\odot \sim 0.02$) and increases slightly with mass. We take the first point where $dL/dt > 0$ to be the starting point of the evolutionary tracks. For $M \lesssim 15 M_\odot$, luminosity increases immediately, starting at $Z = 0$.

3. RESULTS FOR STELLAR EVOLUTION AT ZERO METALLICITY

3.1. Evolutionary Tracks

We have calculated evolving stellar models for $M = 1$ – $100 M_\odot$ from the ZAMS past the onset of helium burning. These tracks appear in Figure 1, which shows the tracks compared with a similar Population II ($Z = 0.001$) ZAMS for $M = 1$ – $100 M_\odot$ from Schaller et al. (1992). The large gain in effective temperature at $Z = 0$ is readily apparent in this figure.

We have done a simple comparison of our tracks to the similar work of Marigo et al. (2001) and Siess et al. (2002). Our tracks match theirs on the H-burning MS and for He-burning at $M > 10 M_\odot$. We find effective temperatures that are 3%–5% lower and H-burning lifetimes that are 1%–12%

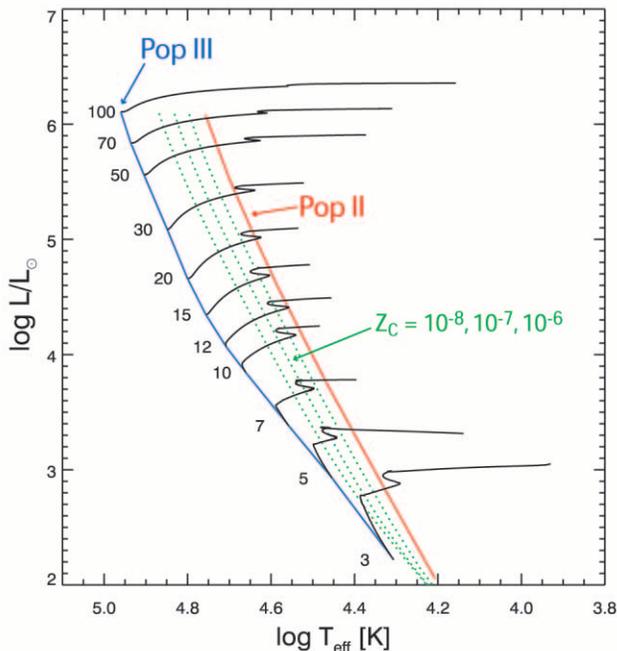


FIG. 1.—The $Z = 0$ stellar evolution tracks from 3 to $100 M_\odot$. The $1 M_\odot$ track has been omitted for clarity. These tracks terminate when the convective core expands to reach the H-rich region during core He burning. To illustrate the stark differences between Population III and their metal-enriched counterparts, we mark in red the ZAMS for Population II ($Z = 0.001$) from Schaller et al. (1992). The Population III ZAMS is marked in blue, and the ZAMSs for $Z_C = 10^{-8}$, 10^{-7} , and 10^{-6} are plotted with dashed curves, with metallicity increasing to the right.

TABLE 2
POPULATION III H-BURNING
LIFETIMES

Mass (M_\odot)	τ_{H} (yr)
1.....	5.87×10^9
3.....	2.01×10^8
5.....	6.03×10^7
7.....	3.11×10^7
10.....	1.76×10^7
12.....	1.40×10^7
15.....	1.11×10^7
20.....	8.47×10^6
30.....	5.86×10^6
50.....	4.09×10^6
70.....	3.42×10^6
100.....	2.93×10^6

shorter (Table 2). These differences illustrate that individual choices for parameters, numerical techniques, and the specific implementation of physical processes (e.g., convection) or boundary conditions can create small differences in the final results. Our models diverge from Marigo et al. (2001) in the details of evolution in the H-R diagram during the He-burning phases for $M \leq 10 M_\odot$, perhaps owing to differences in our treatment of convection. Since we are concerned with the spectral evolution for the first 10 Myr, which is dominated by the H-burning MS of massive stars, we have not attempted to resolve these differences here. Simple testing indicates that the time evolution of ionizing photons and broadband spectra are similar between the two sets of tracks, except for a roughly 30% smaller He II ionizing photon production from our models, due to the slightly cooler effective temperatures and the sensitivity of 4 ryd radiation to T_{eff} . We note also the studies of Population III stars over a broad mass range by Schaerer (2002) and Bromm, Kudritzki, & Loeb (2001), who published static models of zero-age Population III stars from 100 to $1000 M_\odot$.

Our tracks do not complete core He burning because of a peculiar evolutionary feature noted also by Marigo et al. (2001). During core He burning, the core begins to expand and at some point reaches the boundary of the H-depleted region. At this point, the core H abundance rises from 0 to $\sim 10^{-4}$ from one discrete time step to the next, owing to the assumption of instantaneous complete mixing in the convective core. This assumption is clearly inadequate here, since the added H is probably burned quickly at the edge of the convective region in a continuously evolving star. An ad hoc treatment of this phenomenon that forbids expansion of the convective core past this boundary is unlikely to represent accurately the true behavior of the stellar interior in this stage. Since this evolutionary phase represents a small fraction of He burning, which is itself $\lesssim 10\%$ of the stellar lifetime, we terminate our tracks when the boundary of the convective core reaches the H-rich region.

Figure 1 shows that, for $M \gtrsim 12 M_\odot$, the stars evolve uniformly toward lower T_{eff} and higher luminosity. These stars are powered by core CNO burning, catalyzed by the small C fraction built up there by triple- α burning. For $M \lesssim 12 M_\odot$, the stars start on the ZAMS with $Z = 0$ and first evolve to higher T_{eff} via p - p burning while the core C abundance rises. When the ^{12}C mass fraction reaches $Z_C \sim 10^{-9}$, the star is

powered primarily by CNO burning and begins to evolve to lower T_{eff} values.

To assess the differences in ionizing photon emission and some observational characteristics between Population III and Population II, we have constructed ZAMSs for initial carbon abundances of $Z_C = 10^{-8}$, 10^{-7} , and 10^{-6} . These appear in Figure 1. These unevolved populations may represent the intermediate stellar generations between Population III and Population II. We construct model spectra for these populations as for the Population III evolving cluster (see § 4).

We have neglected mass loss in calculating our stellar evolution models and spectra. Mass loss is believed to be driven in massive stars by radiation pressure on blanketed metal lines in the outer reaches of the stellar atmospheres. In the metal-free case, this pressure is likely to be minimal (Kudritzki 2000), and the constant-mass assumption is a good approximation. Marigo et al. (2001) used a simple test of Eddington luminosity to argue that massive Population III stars may experience modest mass loss late in their He-burning phases. Even if this occurs, it is unlikely to affect the composite spectra that we produce here and can be safely ignored. The uncertain role of mass loss can be viewed as one of the major systematic uncertainties remaining in the study of metal-free stars. If it is strong, MS mass loss drives the star to hotter T_{eff} values on the MS; these effects can be seen in the tracks of El Eid et al. (1983) and Klapp (1983).

3.2. Why Population III Stars are Hot

The distinctive spectral appearance of Population III stars is due directly to their unusually hot surfaces (TS). These stars are roughly 2 times hotter and 5 times smaller than their Population II counterparts of the same mass. There are two main reasons for these differences, both related directly to the chemical composition. First, nuclear burning in a metal-poor isotopic mix must occur at higher temperatures to derive the same energy per unit mass as at higher metallicity. The change in nuclear energy production rates ϵ is illustrated in Figure 2, with some characteristic parameters. The highest curve (*green curve*) is ϵ_{CNO} for a metallicity of $Z_C = 10^{-6}$. At this metallicity, CNO burning at a typical core temperature of 4×10^7 K can achieve the typical ϵ for a particular location in the stellar interior. If this star is restricted to p - p burning (*black curve*) or to CNO burning at a very low C abundance (*red curve*), it must do so at temperatures above 10^8 K. These latter two cases illustrate Population III nuclear burning. This effect accounts for most of the large changes in stellar structure as metallicity approaches $Z = 0$. Second, and less important to the structure of Population III stars, the opacity of stellar material is reduced at low metallicity, permitting steeper temperature gradients and more compact configurations at the same mass (see Fig. 3).

The effects of reduced opacity can be seen in Figure 3, which shows the run of six important physical quantities in Population II (*dashed curves*) and Population III stars (*solid curves*) of $15 M_{\odot}$. Population III stars are more centrally condensed, with hotter, denser cores and hotter surfaces. Throughout 99% of the stellar material, the opacity (Fig. 3e) has a weak dependence on metallicity because it is dominated by scattering off electrons donated primarily by H. In the outer envelope, which comprises 1% of the stellar mass, the opacity is provided by the bound-free and bound-bound

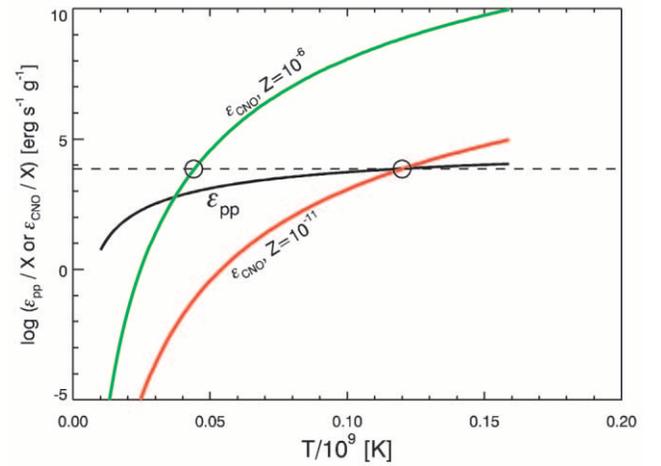


FIG. 2.—Energy generation rates ϵ illustrating why metal-free stars must attain unusually high temperatures to support their mass against gravity. The black curve shows the p - p rate for a typical H mass fraction $X = 0.7$. The green and red curves show the CNO rate for a fiducial low metallicity and for Population III, respectively. The horizontal dashed line marks a typical (but arbitrary) rate for the interior of a massive ($M > 20 M_{\odot}$) star. The circles mark the points where the respective burning processes produce the desired total ϵ . See text for discussion.

transitions of H I, He I, and He II. Here the opacity differs substantially between Population II and Population III, but it has little effect on the overall stellar structure.

4. MODEL ATMOSPHERES AND SYNTHETIC SPECTRA

4.1. Model Atmospheres

The model atmosphere grid was calculated using the publicly available TLUSTY code (Hubeny & Lanz 1995). These models are parameterized by the effective temperature at the stellar photosphere (T_{eff}) and the surface gravity (g). Model spectra at the grid points are calculated for $\lambda = 1\text{--}10,000$ Å by the companion program SYNPLLOT. For high accuracy and coverage, we choose $\Delta T_{\text{eff}} = 500$ K and $\Delta(\log g) = 0.05$. Figure 4 shows the range of the model grid with the evolutionary tracks superposed. At low temperatures ($T_{\text{eff}} < 10,000$ K), the model atmospheres are allowed to have small convective regions to promote convergence. This modification has little effect on their emergent spectra. Some stars develop convective envelopes during late H burning, but these regions lie far from the nuclear burning core and thus always maintain their metal-free composition.

4.2. Population Synthesis

To produce model spectra and calculate the cosmological effects of Population III stars, we must create synthetic clusters from the discrete set of tracks presented above. We desire smoothly varying quantities such as ionizing photon flux and broadband colors, so we must convert the tracks at discrete masses into more continuous variables.

Constructing a complete synthetic population from evolutionary tracks requires careful calculation of the tracks for masses that are not explicitly computed. For example, if one has tracks at 100 , 80 , and $60 M_{\odot}$, extrapolating the behavior of the $90 M_{\odot}$ track from 60 and $80 M_{\odot}$ leads to large errors, and interpolation of the 100 and $80 M_{\odot}$ tracks is not accurate after the last time step on the more massive

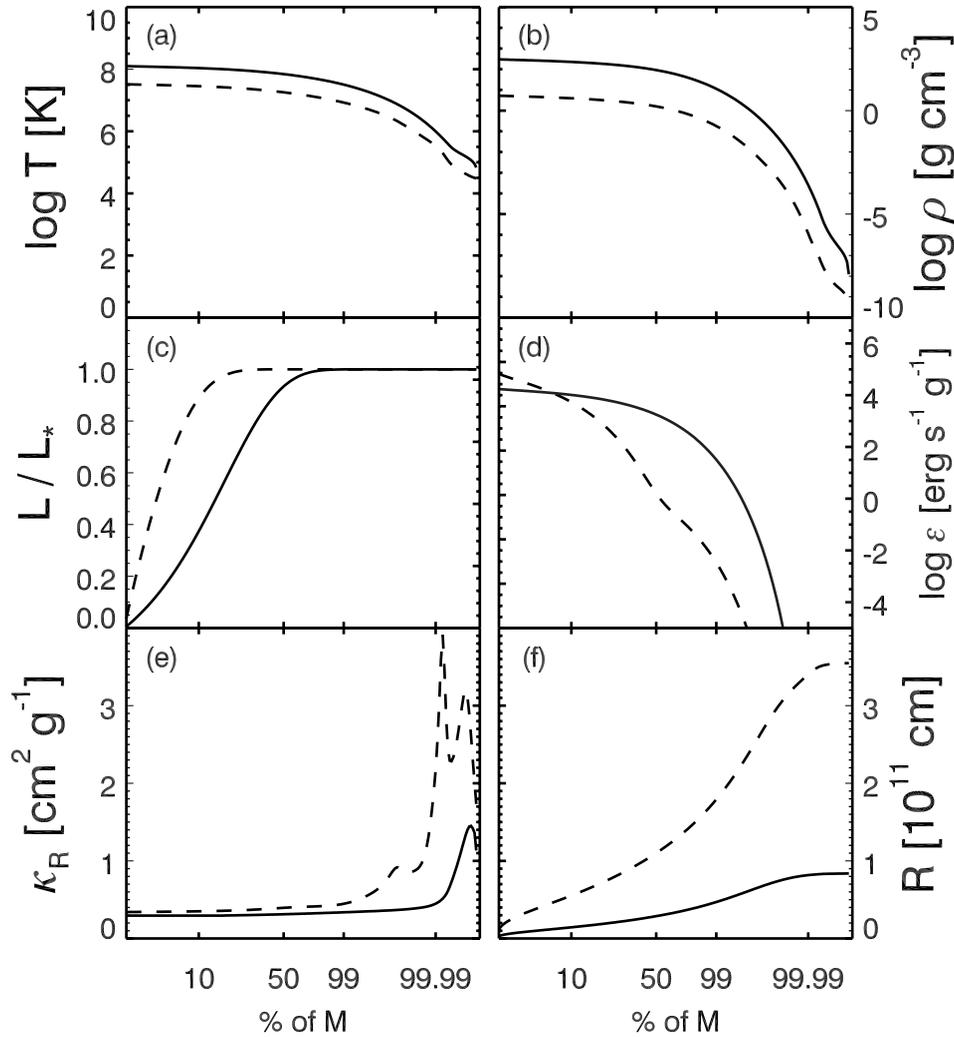


FIG. 3.—Comparison of the interior structure of a Population III star (*solid curves*) with a Population II star (*dashed curves*), both of $M_* = 15 M_\odot$. The Population III star is smaller, hotter, and denser. This comparison demonstrates that the gain in T_{eff} arises not from changes in the opacity with metallicity, which differs substantially only in the outer 1% of the total mass, but rather from changes in the available fuel source in the stellar core. See § 3.2 for discussion.

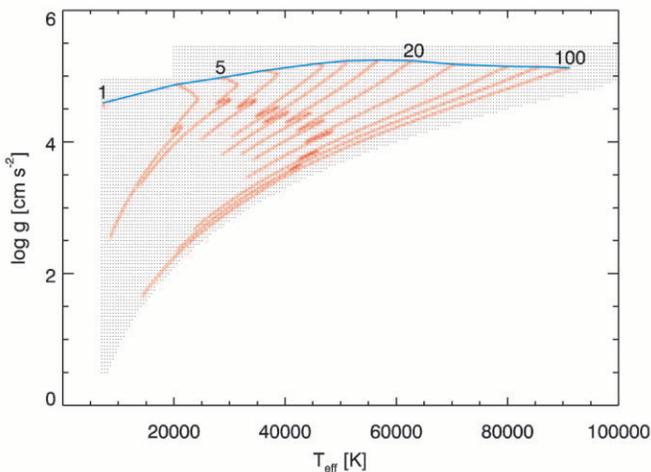


FIG. 4.—Parameter grid (T_{eff} , $\log g$) of the TLUSTY non-LTE model atmosphere grid with the Population III tracks superposed. The grid boundaries were chosen especially to cover the tracks with a minimum of extraneous points. The grid spacing is $\Delta T_{\text{eff}} = 500$ K and $\Delta(\log g) = 0.05$. Only the first 10^9 yr of the $1 M_\odot$ track is accommodated by the grid. The Population III ZAMS is marked in blue.

track. We cannot simply use the explicitly calculated tracks to represent broad mass bins ($\Delta M \sim 20 M_\odot$), because this leads to sharp discontinuities in the calculated spectra and derived quantities. We use small mass bins of $1 M_\odot$, which are preferred for adequate accuracy and smooth behavior of the results. To resolve these issues we use a “lifetime-proportional” interpolation scheme to create synthetic evolutionary tracks for the intermediate masses between the explicitly calculated tracks.

The evolutionary tracks at the explicitly calculated masses are followed beyond the onset of helium burning, a phenomenon that occurs at all masses and occupies $\lesssim 10\%$ of the stellar lifetime. Since later stages (C burning, Si burning) are even shorter and contribute little to the spectral evolution of a cluster, we assume that He burning is the last evolutionary phase and that its end terminates the star’s life. We also rely on the fact that the changes in the shape and time dependence of the tracks vary smoothly with mass. We can use these important features to construct interpolated tracks at intermediate masses.

Taking the run of $T_{\text{eff}}(t)$ and $L(t)$ to be a track, we first remap the explicit tracks to a new time coordinate: the fraction of the stellar lifetime achieved to that point, ranging

from 0 to 1. Then, a vector of lifetimes for each mass bin is interpolated from the lifetimes of explicit tracks. Finally, we interpolate intermediate tracks between the explicit tracks at the same fraction of their lifetimes. For instance, to find (L, T_{eff}) for a $90 M_{\odot}$ star at 70% of its interpolated lifetime, we interpolate in mass between the 80 and $100 M_{\odot}$ explicit tracks at 70% of their lifetimes. This scheme minimizes interpolation errors and generates smoothly varying model spectra and derived quantities.

Using the interpolated tracks and model atmospheres, we construct synthetic stellar populations by specifying a stellar initial mass function (IMF) and adding stellar spectra together. Because we lack detailed knowledge of the IMF for $Z = 0$ star formation, we assume that the IMF is a power law, $dN/dM \propto M^{-\alpha}$, with the Salpeter slope, $\alpha = 2.35$. To populate the IMF, we construct 99 mass bins and populate them with the number of stars of that mass, normalized to a total mass of $10^6 M_{\odot}$, which is standard practice. The mass bins are $1 M_{\odot}$ wide and range from 1 to $100 M_{\odot}$. Model spectra from the atmosphere grid are assigned to the bins and summed over all masses. The product of this process is a composite spectrum for the whole cluster. Schaerer (2002) and Bromm et al. (2001) performed similar calculations for top-heavy IMFs ranging up to $1000 M_{\odot}$.

We adopt a model of Population II starburst clusters as a reference spectrum for comparison with our evolving Population III models. In the present paper, this reference model represents a baseline spectrum for comparison to the $Z = 0$ models. In Paper II, we use this spectrum as a proxy for the “second generation” of stars incorporating the metals produced by the first. For the reference population, we use a model (with $Z = 0.001$, $M = 10^6 M_{\odot}$ in stars and a Salpeter IMF) that was custom generated by the Starburst99 population synthesis code (Version 3.1; Leitherer et al. 1999). Because this suite of models is widely used in the community, we compare all our model spectra and observational predictions to a Population II model generated by their code. We note further that we do not include Wolf-Rayet (W-R) stars in these comparison models to make our comparisons as direct as possible. W-R stars are known to produce large amounts of He II ionizing photons in the local universe (Schaerer & Vacca 1998), but their presence at high redshift ($z > 5$) and low metallicity is regarded as unlikely since the W-R phenomenon requires metal-line opacity to drive strong stellar mass loss. Recent studies of stellar mass loss at low metallicity (Kudritzki 2002) have shown that its importance is minimized at extremely low metallicity, with a factor of 10 drop in mass-loss rate from $Z = 10^{-3}$ to 10^{-4} . Nonetheless, we include in our comparisons below some results for Population II with W-R stars.

For modeling cosmological reionization, radiative feedback, and the observational signatures of metal-free stars, we make use of the important quantities Q_i , which express the ionizing photon output per second in the H I, He I, and He II continua ($i = 0, 1, 2$ and $h\nu_i = 13.6, 24.6,$ and 54.4 eV, respectively):

$$Q_i = 4\pi R_*^2 \int_{\nu_i}^{\infty} \frac{F_{\nu}}{h\nu} d\nu, \quad (8)$$

where F_{ν} is the spectrum of the star in units of $\text{ergs cm}^{-2} \text{s}^{-1} \text{Hz}^{-1}$. The time-dependent values of Q_0 and Q_2 are

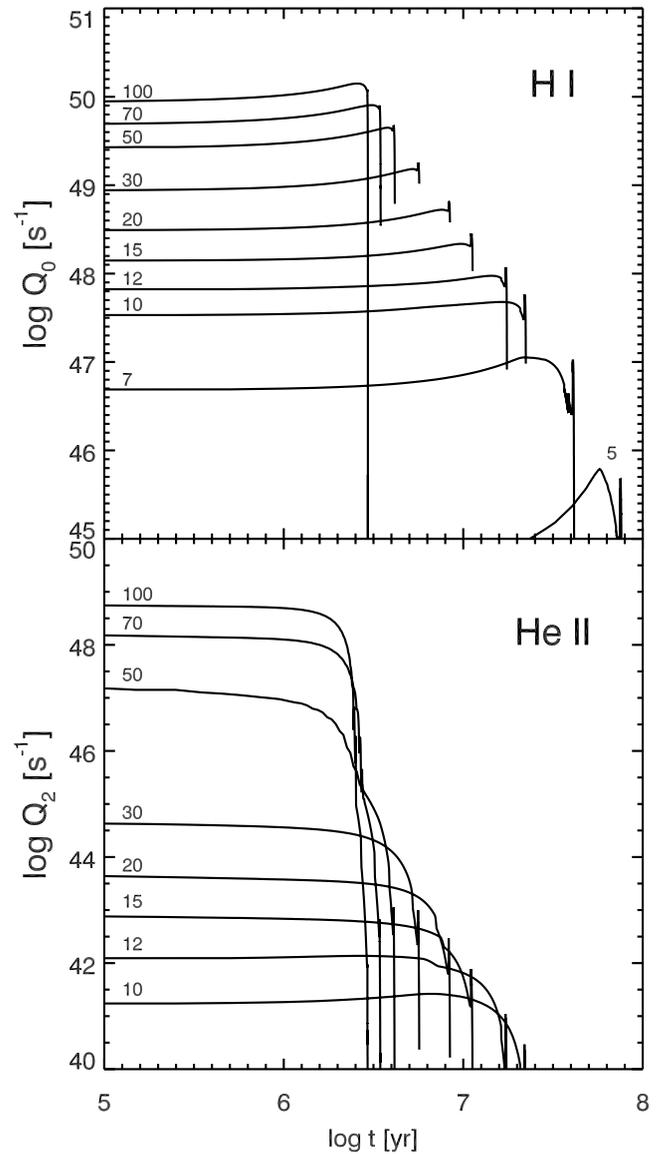


FIG. 5.—Ionizing photon production rates Q_0 (1 ryd; top) and Q_2 (4 ryd; bottom) for individual stars from the tracks in Fig. 1. The tracks are labeled with their stellar masses in units of solar mass. High-mass, zero-metallicity stars are strong sources of H I and He II ionization. In particular, they emit 10^5 times more radiation in the He II continuum than Population II stars.

displayed for the individual stars in Figure 5 to show the behavior with mass.

Once the individual Q_i have been weighted by the IMF, we arrive at the cluster Q_i seen in Figure 6. There are several features of this figure that deserve comment. First, note the starting (ZAMS) values of Q_0 and Q_2 for the Population III and comparison Population II clusters. The Population III cluster has a Q_0 that is 60% larger than the Population II cluster. Second, the ionization of the Population III cluster persists at a high level far longer than the Population II cluster, owing to the strong gain in ionization for $M = 10$ – $25 M_{\odot}$ stars. The Population III cluster takes 10 Myr to decline to 1/10 its ZAMS value in Q_0 , twice as long as Population II. Finally, note the strong ionization above the He II edge at 4 ryd. At the Population III ZAMS, we find that $Q_2/Q_0 = 0.02$, with a lifetime of 2.5 Myr, characteristic of $M \geq 40$ – $50 M_{\odot}$ stars (see Fig. 5). This form of the output

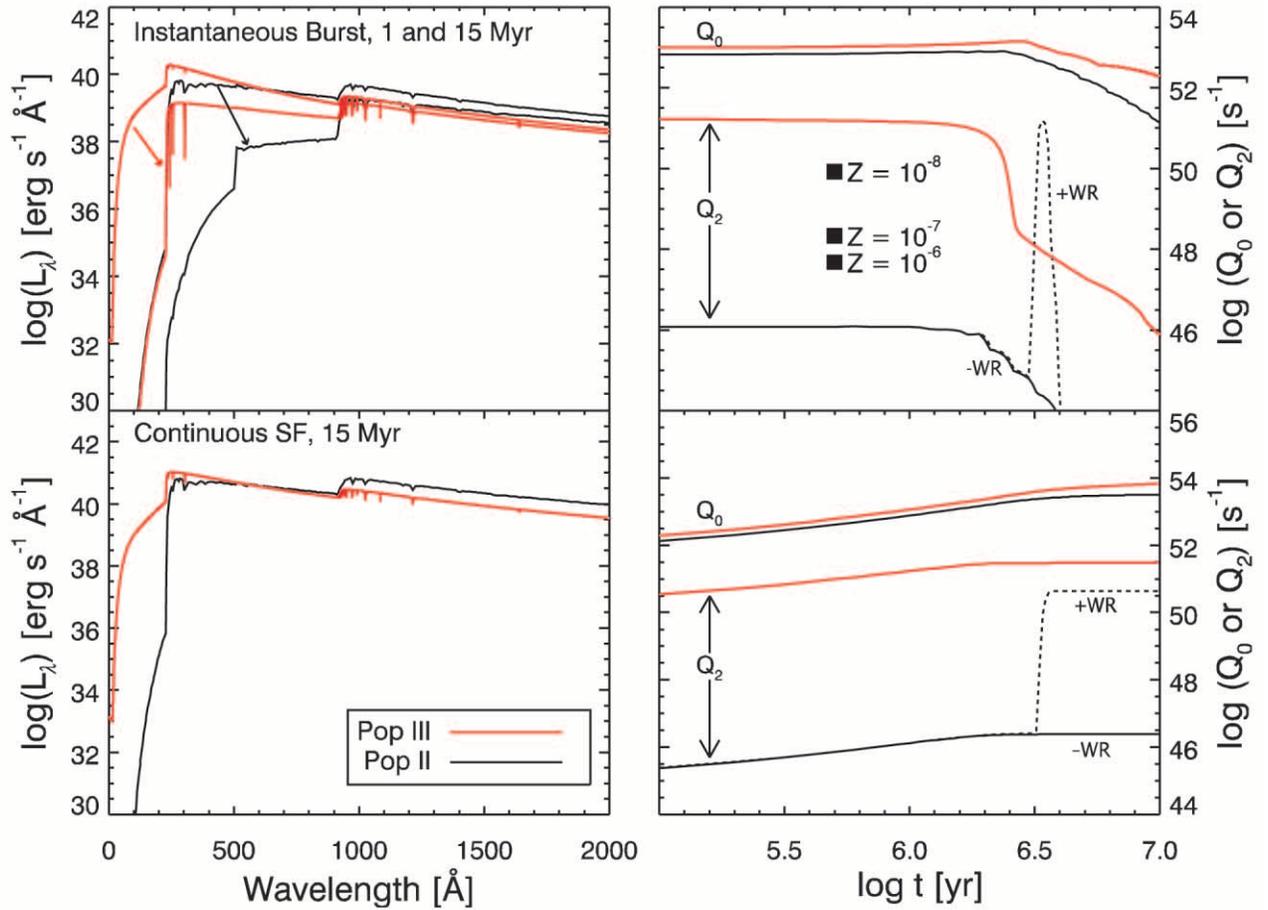


FIG. 6.—*Top left:* Composite spectra for Population II (*black curve*) and Population III (*red curve*) clusters ($M = 10^6 M_\odot$ in stars) at 1 and 15 Myr. At 15 Myr the Population II spectrum has faded in H I ionization, but the Population III cluster is still a significant source of H I ionization owing to the presence of stars with $M = 10\text{--}15 M_\odot$ and $T_{\text{eff}} > 50,000$ K. No nebular emission is included here. *Bottom left:* Example spectra for the continuous star formation case at 15 Myr. *Top and bottom right:* Cluster Q_0 and Q_2 corresponding to the instantaneous/continuous cases at left. H I ionization from the Population III cluster is 60% stronger than that from Population II. Population III emits 10^5 times more He II ionizing photons than Population II with the same IMF and total mass, excluding W-R stars. If W-R stars are included, this gain in He II ionization is smaller but still substantial. This key result has potentially large effects on the IGM (see Paper II). In the top right panel we mark with filled squares the instantaneous values of Q_2 from the ZAMSs with $Z_C = 10^{-8}$, 10^{-7} , and 10^{-6} (plotted at an arbitrary time). These points show the sharp decline in He II ionizing photon production when small abundances of ^{12}C are included. The right panels also show the Starburst99 results with W-R stars (*dotted curves*), for comparison. See § 4.2 for discussion.

from our models can be placed directly into cosmological calculations of IGM reionization that incorporate realistic models of evolving spectra from Population III stars (see Paper II).

For the ZAMS spectra at $Z_C = 10^{-8}$, 10^{-7} , and 10^{-6} , we find $Q_0 = (8.9, 8.4, 7.9) \times 10^{52} \text{ s}^{-1}$, respectively, intermediate between the Population II and Population III values. There is a sharp decrease in their Q_2 values, however, with $Q_2 = 8.4 \times 10^{49}$, 2.1×10^{48} , and $4.8 \times 10^{47} \text{ s}^{-1}$, respectively. These quantities are compared with the evolving Q_2 values for Population III and Population II in Figure 6.

In Paper II, we examine the effects of Population III clusters on the reionization of the IGM. These tests use a quantity η , which expresses the ratio of He II to H I column densities in an ionized IGM (Fardal, Giroux, & Shull 1998). This ratio can be predicted for a specific distribution of IGM clouds, given an input source spectrum, which is assigned an intrinsic η (prior to IGM filtering) based on the shape of its spectrum above 1 ryd. A composite QSO (Zheng et al. 1997; Telfer et al. 2002) with a power-law spectrum,

$F_\nu \sim \nu^{-\alpha}$, and $\alpha = 1.8$ has an intrinsic $\eta = 20$. At the ZAMS, the Population III cluster has $\eta = 10$, harder than the mean QSO and similar to a power-law spectrum with $\alpha = 1.3$. For the low-metallicity ZAMS at $Z = 10^{-8}$, 10^{-7} , and 10^{-6} , we find intrinsic $\eta = 33$, 80, and 180, respectively. We make use of this ratio in Paper II to explore the fate of IGM regions ionized by Population III stars.

5. OBSERVATIONAL SIGNATURES

5.1. General Approach

Observational searches for high-redshift galaxies work best when equipped with diagnostic tools that provide detailed relationships between observable signatures and underlying physical phenomena. To this end, we use our models of Population III clusters to derive predictions for the observed spectral energy distributions (SEDs) of the first stars. We consider two classes of observation: the first seeks the nebular emission-line signatures of Population III stars, and the second exploits their broadband SED and colors.

We take up the nebular emission first because its presence affects the broadband signatures as well.

5.2. Emission-Line Diagnostics

Emission-line signatures are proving to be the primary means of finding galaxies at redshifts beyond $z = 5$. Deep, blind searches for Ly α emitters have turned up large populations of galaxies at these epochs. These surveys either use blind multislit spectrographic exposures to detect single emission lines over a broad redshift range (Hu et al. 1999) or take narrowband images to probe deeply at selected redshifts (Rhoads et al. 2000).

TGS derived a relationship for the observed He II emission-line flux from metal-free stars as a function of the SFR. With the evolved spectra of metal-free stars we can update these earlier calculations. We assume that few He II ionizing photons escape the galaxies where they are produced ($f_{\text{esc}} = 0.025$; see Paper II) and that the sources themselves contain no dust. We assume case B recombination at 20,000 K, corresponding to a higher nebular temperature in low-metallicity gas. As in TGS, we adopt $j_{4686}/j_{\text{H}\alpha} = 0.66$, $j_{1640}/j_{\text{H}\alpha} = 4.6$, and $j_{3203}/j_{\text{H}\alpha} = 0.28$ (Seaton 1978). TGS used an ad hoc scale factor f_{evol} to account for stellar evolution effects, which reduce the He II ionizing flux over time as high-mass stars evolve off the MS. We need not make this approximation here, since we have directly calculated evolving spectra of metal-free stars. For both continuous and instantaneous star formation, we assume that $L_{\text{H}\alpha} = 1.27 \times 10^{-12} Q_0$ ergs s $^{-1}$ (i.e., there are 0.424 H α photons emitted for each H-ionizing photon). Similar coefficients are derived for the other H I lines from case B recombination theory (Osterbrock 1989). Note that Leitherer et al. (1999) use a coefficient of 1.36×10^{-12} , corresponding to nebular gas at 10^4 K. We then multiply the H α luminosity by the emissivity ratio $j_i/j_{\text{H}\alpha}$ to derive L_i in a He II line. In Figure 7 we plot the predicted line fluxes for Ly α and He II $\lambda 1640$ as a function of redshift.

For clusters with significant ionizing photon production, we must also account for continuous nebular emission from the H I two-photon process and direct recombinations to upper levels of H I, He I, and He II. We construct a wavelength-dependent emission spectrum, in units of ergs s $^{-1}$ \AA^{-1} :

$$F_{\lambda}^{\text{cont}} = \frac{c\gamma_i}{\lambda^2\alpha_{\text{H}}^{\text{B}}} (1 - f_{\text{esc}}) Q_0, \quad (9)$$

where the term γ_i , in units of ergs cm 3 s $^{-1}$ Hz $^{-1}$, incorporates contributions from H I, He I, He II, and the H I two-photon continuum and $\alpha_{\text{H}}^{\text{B}}$ is the case B recombination rate coefficient for hydrogen. The emission coefficients for an electron temperature of 20,000 K are taken from Aller (1984, Table 4.9). This additional emission has a large effect on the broadband SED of the synthetic clusters, shifting energy from the Lyman continuum to the rest-frame optical bands. This effect is included in the broadband color and equivalent width predictions presented below.

In general, we can parameterize the escape of radiation with the quantity f_{esc} , the fraction of ionizing photons that escape from the host halo into the IGM. To be consistent with the typical reionization models presented in Paper II, we uniformly assume $f_{\text{esc}} = 0.05$ for 1–4 ryd and $f_{\text{esc}} = 0.025$ for $h\nu > 4$ ryd (Dove, Shull, & Ferrara 2000; Tumlinson et al. 1999), leaving most ionizing photons at all wavelengths

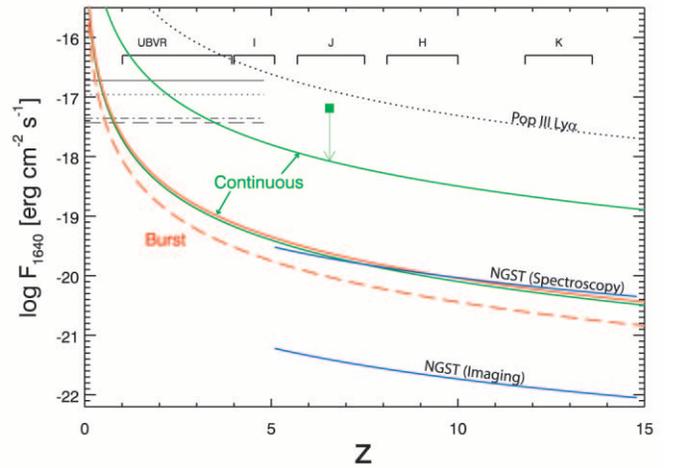


FIG. 7.—Predicted fluxes for Ly α and He II $\lambda 1640$ from the Population III cluster. Ly α emission may serve as signposts pointing the way to metal-free star formation, which could then be confirmed by the detection of He II lines. The green square marks the observed Ly α flux detected from the $z = 6.56$ galaxy found by Hu et al. (2002), and the vertical arrow connects this point to the predicted He II flux curve if this galaxy harbors a continuous Population III starburst with SFR = $40 M_{\odot} \text{ yr}^{-1}$. The lower green curve corresponds to SFR = $1 M_{\odot} \text{ yr}^{-1}$. The red curves mark the He II $\lambda 1640$ flux for an instantaneous burst at 0 and 1 Myr (solid and dashed curves, respectively). The blue curves plot the imaging and spectroscopic sensitivities for JWST. At the top are wavelength ranges of the common photometric bandpasses. The horizontal lines extending to $z = 5$ represent the sensitivity limits of recent emission-line surveys of galaxies: from top to bottom, Rhoads et al. (2000), Hu et al. (1999; imaging data), Stern et al. (1999), and Hu et al. (1999; spectroscopic data).

to excite observable nebular emission. We return to the issue of the escape fraction in Paper II, where we discuss justifications for different escape fractions at 1 and 4 ryd.

In Figure 7 we show the flux of the $\lambda 1640$ and Ly α lines from the Population III cluster. The green curves represent continuous star formation at 1 and $40 M_{\odot} \text{ yr}^{-1}$. The higher rate corresponds to the SFR inferred by Hu et al. (2002) for a Ly α emitter discovered at $z = 6.56$ (the Ly α flux marked with the green square). The red curves represent the He II $\lambda 1640$ flux from an instantaneous burst of $M = 10^6 M_{\odot}$ at $t = 0$ (solid curve) and 1 Myr after the burst (dashed curve). Figure 7 also shows the imaging and spectroscopic sensitivity requirements for the planned 6 m aperture *James Webb Space Telescope* (JWST). These assumed limits are from the JWST design specifications and are 2 nJy for low-resolution ($R = 5$) imaging and 100 nJy for spectroscopy ($R = 1000$), and both assume S/N = 10 and a 100,000 s exposure time. These comparisons show that the planned JWST will easily discover this signature of the first stars.

In case B recombination theory, roughly two-thirds of all H recombinations result in the emission of an Ly α photon (Spitzer 1978), and these are known to escape from galaxies and serve as a strong signature of high-redshift star formation (Hu et al. 1999; Rhoads et al. 2000). For the Population III cluster presented here, the Ly α line is intrinsically stronger than in Population II. Figure 8 shows the predictions for the time-dependent Ly α equivalent width $W_{\text{Ly}\alpha}$ for Population II and Population III assuming case B recombination theory and $T = 20,000$ K. These curves are technically upper limits since the Population II cluster may suffer dust attenuation and the Population III Ly α line may be absorbed by neutral H in the pre-reionization IGM

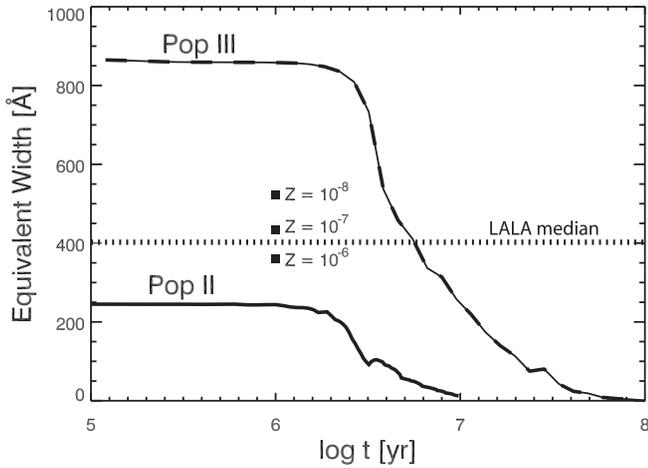


FIG. 8.—Predicted equivalent width W_λ in the Ly α emission line for the Population III synthetic cluster (a Salpeter IMF up to $100 M_\odot$). The Population III equivalent width is unusually large, owing to the harder Population III spectra. The horizontal dotted line marks 400 Å, the median equivalent width for 157 galaxies in the LALA sample (Malhotra & Rhoads 2002). The filled squares mark the instantaneous $W_{\text{Ly}\alpha}$ from the low-metallicity ZAMS, placed arbitrarily at $t = 10^6$ yr for clarity. Population III clusters may be identified by their large Ly α equivalent widths before their continuum spectra or broadband colors are measured. However, we note that at Population II metallicities, a top-heavy IMF or one with stars more massive than $100 M_\odot$ can mimic this effect.

(see Loeb & Rybicki 1999). Equivalent width has the additional benefit of being independent of the total cluster mass for an instantaneous burst. Thus, we confirm the suggestion of other studies (TS; Bromm et al. 2001; Schaerer 2002) that the Ly α signatures of the first stars will be unusually strong.⁴ The Large Area Ly α Survey (LALA; Malhotra & Rhoads 2002) found a median $W_{\text{Ly}\alpha}$ of 400 Å in a sample of 157 galaxies with $z = 4.5$. More than 60% of their detections have $W_{\text{Ly}\alpha} > 240$ Å, the approximate maximum attainable by Population II clusters (Charlot & Fall 1993).

We also plot in Figure 8 the instantaneous equivalent widths $W_{\text{Ly}\alpha} = 520, 430,$ and 360 Å for the $Z_C = 10^{-8}, 10^{-7},$ and 10^{-6} ZAMS spectra, respectively. The points for $Z_C = 10^{-7}$ and 10^{-6} straddle the median LALA equivalent width, indicating that this sample of Ly α emitters may have a low metallicity. These results suggest that populations of metal-free stars, or perhaps stars of the first metal-enriched generation, have already been found. In future ground- and space-based searches beyond the optical bands, Ly α emission will probably serve as the “signpost” of metal-free galaxies. This result could then be confirmed through follow-up detection of the He II lines, which are weaker than Ly α and unique to the metal-free regime.

The sensitive dependence of the Ly α equivalent width on metallicity (and IMF; see Bromm et al. 2001) suggests that the metallicity evolution of star formation may appear

⁴ We note here that our Ly α predictions lie well below those of Schaerer (2002) for similar conditions. We use an average flux in the 1200–1250 Å range to represent the continuum flux. This method approximates the observed quantities in surveys such as LALA. If intrinsic Ly α absorption is present, our prediction is an overestimate. We can reproduce Schaerer’s prediction only by assuming the continuum flux is represented by the core of the model Ly α absorption line.

in the observed Ly α lines at high redshift. If there is a well-defined Population III epoch, characterized by a rapid transition to $Z_C = 10^{-6}$ or higher as the second stellar generation forms from gas enriched by the first, there could be a bimodal distribution of Ly α equivalent widths with redshift. The peak may appear near the ~ 1000 Å prediction here, nearer the ~ 3000 Å predicted for the high-mass IMF adopted by Bromm et al. (2001), or somewhere in between. Thus, Ly α itself may serve as a diagnostic of the evolution of metals in galaxies, up to at least the redshift of H reionization, where the damping wing of the Gunn-Peterson absorption in the IGM plays an uncertain role in obscuring (Miralda-Escudé 1998) or transmitting (Loeb & Rybicki 1999; Haiman 2002) the galactic Ly α line.

5.3. Broadband Spectra and Colors

Broadband filter search techniques have proven successful in finding high-redshift galaxies. In particular, the “Lyman break” technique exploits the fact that the IGM absorbs the spectrum of a galaxy shortward of $1216(1+z)$ Å. This powerful technique has uncovered thousands of galaxies at $z > 3$ (Steidel et al. 1999).

The high effective temperatures of Population III stars make their integrated cluster spectra intrinsically bluer than their Population II counterparts. However, the most distinctive feature of the Population III spectrum is the large gain in H I and He II ionizing photons, which unfortunately are emitted in the spectral region that is heavily absorbed by the IGM. To account for this obscuration, we adopt a stochastic model of the IGM based on the work of Fardal et al. (1998), who used Monte Carlo techniques to construct model sight lines containing realistic distributions of H I Ly α forest, Lyman limit, and damped Ly α clouds. From a sample of 1000 sight lines we construct a mean transmission function for three redshifts, $z = 4, 5,$ and 6 . This realistic treatment of the IGM leaves approximately 30% of the emergent flux between $912(1+z)$ and $1216(1+z)$ Å. This mean transmission function is included in the broadband signatures derived here.

In Figure 9, we illustrate Population II and Population III spectra at high redshift. We also plot the ranges of the standard photometric bandpasses. Emission lines are included in these colors and make a modest contribution. Although the Population III spectrum is intrinsically bluer in the accessible spectral regions, the reprocessing of ionizing photons into nebular emission makes up for this difference. Figure 10 illustrates this similarity between Population III and Population II over a range of redshifts and cluster ages ($t = 0$ – 10 Myr). This substantial degeneracy in the broadband colors may make distinguishing Population III more challenging for current and planned optical and IR instruments than was previously appreciated. Confusing color signatures may leave the unusual Population III emission-line signatures as the only means of identifying the first stars.

6. CONCLUSIONS

We have presented evolving spectra of metal-free stellar populations, based on newly calculated evolutionary tracks. We find that the evolution of Population III stars follows the general patterns obeyed at higher metallicity, but with an overall shift to higher temperature. This gain

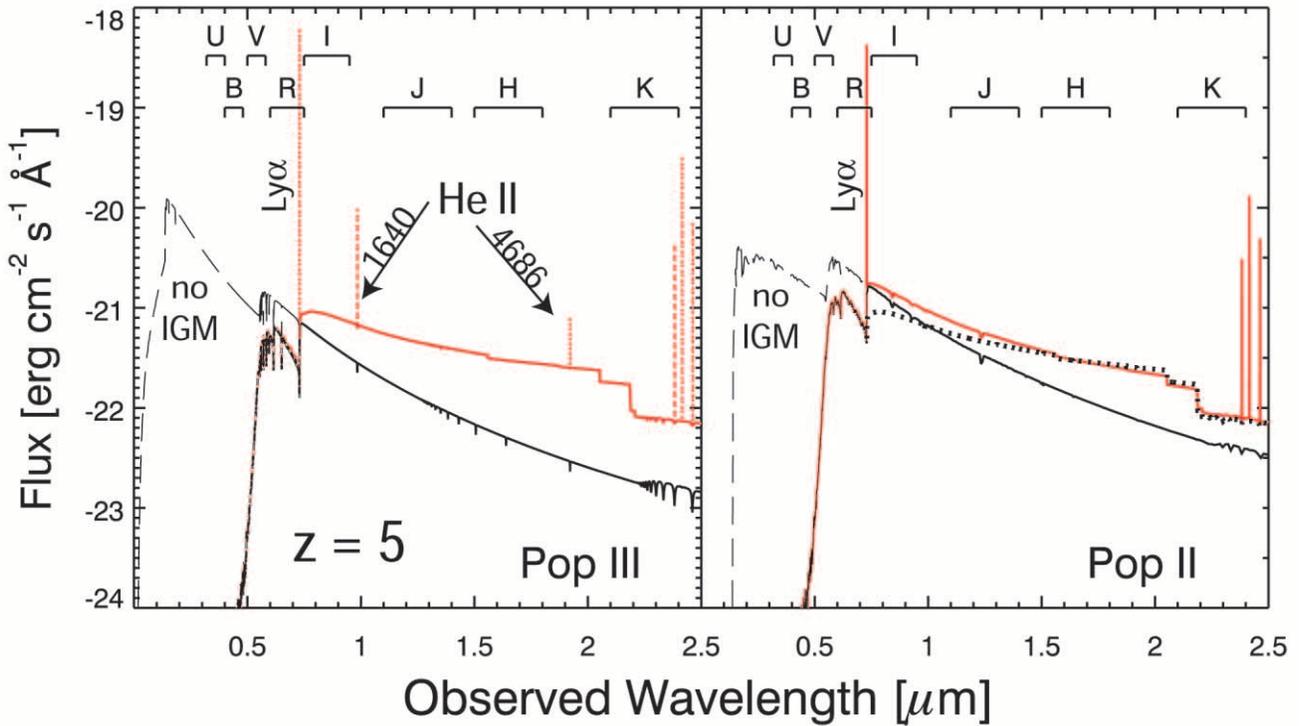


FIG. 9.—Example observed spectra at time $t = 0$ for the Population III (left) and Population II (right) cluster at $z = 5$. Shortward of the rest-frame Lyman limit, the spectra are marked with long-dashed curves. The black curves mark the intrinsic stellar spectra, and the red curves include the nebular continuum emission. Labeled along the top are the common optical and infrared photometric bandpasses. The dotted curve in the right panel repeats the Population III spectrum as left for comparison. The spectral regions where Population III luminosity exceeds that of Population II lie in the range heavily attenuated by the IGM [$\lambda < 1216(1+z)$ Å], which is represented here by a mean transmission for 1000 sight lines simulated by a Monte Carlo model of the IGM (Fardal et al. 1998). Note the absence of the He II recombination lines in the Population II spectrum and the similarity between the observed spectra when nebular continuum emission is included.

in core and surface temperatures is primarily a result of the restricted abundance of ^{12}C in primordial stars. We have used a grid of non-LTE model atmospheres to produce evolving spectra of synthetic Population III clusters. These models have been used to compute the broadband colors and emission-line signatures from metal-free stellar clusters at high redshift.

Our specific conclusions fall into two categories. First, the unique composition of the first stars has the following important effects on their spectra:

1. They produce 60% more H I ionizing photons than their Population II counterparts.
2. They can produce up to 10^5 times more He II ionizing photons than Population II, which may lead to unusual radiative feedback effects on the IGM.

We wish to reemphasize here the uncertainty associated with stellar mass loss in evaluating the latter conclusion. If the “second generation” of stars has $Z \sim 0.0001$ – 0.001 and perhaps significant populations of W-R stars, the gain in He II ionizing photon production for a continuous star formation model including W-R stars is reduced to about a factor of 10 (see Fig. 6). However, since the importance of W-R stars at extremely low metallicities is in doubt, we conclude that Population III stars will be far more efficient than their Population II counterparts at ionizing He II. The consequences of our models for reionization are explored more fully in Paper

II, which uses these model spectra and a semianalytic reionization model to assess the importance of the first stars for full or partial reionization of H I and He II in the IGM.

The peculiar properties of metal-free stars entail unusual observational signatures:

1. They are expected to have unusually high equivalent widths of Ly α and He II recombination lines, the most distinctive signature of Population III stars and their nebular emission. These lines will probably be the primary means of detecting and identifying the first stars.
2. While their intrinsic spectra are significantly bluer than their Population II counterparts, their broadband colors are similar, owing to the reprocessing of Lyman continuum photons into continuous nebular emission.

Recent Ly α surveys suggest that metal-free stars may have already been found. In future ground- and space-based searches beyond the optical bands, Ly α emission will probably serve as the “signpost” of metal-free galaxies. Follow-up detections of the He II lines could confirm this result. The substantial degeneracy in the broadband colors will probably leave emission-line techniques as the best way of distinguishing the first stars.

We are grateful to Mark Giroux for calculating the IGM transfer functions presented in § 4 and to Claus Leitherer and his collaborators for the public distribution of

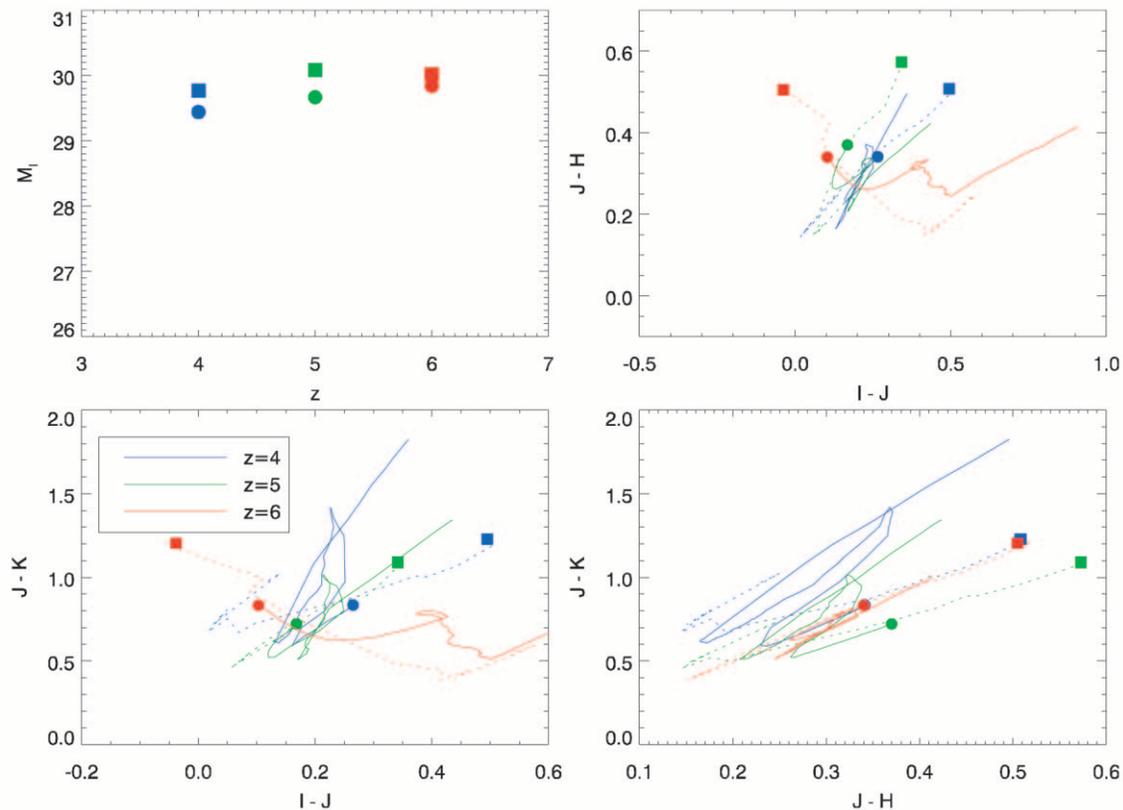


FIG. 10.—Important broadband color diagnostics for the Population III cluster (*dotted curves and filled squares*), compared to Population II (*solid curves and filled circles*) from the ZAMS to 10 Myr in age, at redshifts $z = 4, 5,$ and 6 (*blue, green, and red curves and symbols, respectively*). The tracks start at the filled symbols. Population III clusters are intrinsically fainter and bluer in the infrared bands, but the influence of nebular continuum emission for $t \leq 10$ Myr hides this difference and gives the Population III and Population II clusters similar color signatures. Thus, the recombination lines of H I ($\text{Ly}\alpha$) and He II ($\lambda 1640$) may prove to be the best means of distinguishing Population III stars.

their Starburst99 population synthesis code.⁵ Comments from John Stocke, Nick Gnedin, and Andrew Hamilton improved the manuscript. We happily acknowledge the anonymous referee who prompted us to clarify the discrepancy between our emission-line predictions and those of Schaerer (2002) and who provided other useful comments.

The *JWST* design specifications were taken from the *JWST* Web site in 2002 April.⁶ The OPAL project at Lawrence Livermore National Laboratory is gratefully acknowledged for their opacity data and interpolation codes.⁷ This work was partially supported by an astrophysical theory grant to the University of Colorado by NASA (NAG5-7262).

⁵ Available at <http://www.stsci.edu/science/starburst>.

⁶ See <http://www.stsci.edu/ngst/science/sensitivity>.

⁷ Available at <http://www-phys.llnl.gov/Research/OPAL/opal.html>.

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